

Accretion disc by Roche lobe overflow in the supergiant fast X-ray transient IGR J08408–4503

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Received ...; accepted ...

ABSTRACT

Supergiant fast X-ray transients (SFXTs) are X-ray binary systems with a supergiant companion and likely a neutron star, which show a fast ($\sim 10^3$ s) and high variability with a dynamic range up to 10^{5-6} . Given their extreme properties, they are considered among the most valuable laboratories to test accretion models. Recently, the orbital parameters of a member of this class, IGR J08408–4503, were obtained from optical observations. We used this information, together with X-ray observations from previous publications and new results from X-ray and optical data collected by *INTEGRAL* and presented in this work, to study the accretion mechanisms at work in IGR J08408–4503. We found that the high eccentricity of the compact object orbit and the large size of the donor star imply Roche lobe overflow (RLO) around the periastron. It is also likely that a fraction of the outer layers of the photosphere of the donor star are lost from the Lagrangian point L_2 during the periastron passages. On the basis of these findings, we discuss the flaring variability of IGR J08408–4503 assuming the presence of an accretion disc. We point out that IGR J08408–4503 may not be the only SFXT with an accretion disc fueled by RLO. These findings open a new scenario for accretion mechanisms in SFXTs, since most of them have so far been based on the assumption of spherically symmetric accretion.

Key words. accretion – stars: neutron – X-rays: binaries – X-rays: individuals: IGR J08408–4503

1. Introduction

Supergiant fast X-ray transients (SFXTs) are a subclass of high mass X-ray binaries (HMXBs) composed of a compact object, that is a neutron star (NS) or black hole, accreting material from an OB supergiant star (for recent reviews, see [Romano 2015](#) and [Sidoli 2017](#)). Classical HMXBs with a supergiant companion are X-ray sources with a typical X-ray luminosity of $\sim 10^{35-37}$ erg s⁻¹ and moderate variability, which have a flux dynamic range of about ten. In contrast, SFXTs show sporadic X-ray outbursts with a duration of the order of about a day, composed of short ($\sim 10^2 - 10^3$ s) and bright ($\sim 10^{35-38}$ erg s⁻¹) flares. The low luminosity state is characterised by $L_{\text{low}} \approx 10^{32-33}$ erg s⁻¹, and therefore the flux dynamic range can reach $10^5 - 10^6$. These SFXTs show X-ray spectra reminiscent of those of accreting X-ray pulsars (power law with cut-off in the range 10–30 keV), indicating a likely NS nature for the compact object of the binary system. Several accretion mechanisms have been proposed to explain the X-ray variability of SFXTs. These mechanisms involve the accretion of wind clumps onto the NS (e.g. [in't Zand 2005](#); [Ducci et al. 2009](#); [Chaty 2013](#), and references therein), gating mechanisms ([Grebenev & Sunyaev 2007](#); [Bozzo et al. 2008](#)), subsonic settling accretion regimes (e.g. [Shakura et al. 2014](#)), and formation of transient accretion discs ([Ducci et al. 2010](#)). IGR J08408–4503 has the lowest duty cycle among the SFXTs, with an activity duty cycle of 0.11% (according to the definition of [Paizis & Sidoli 2014](#) based on *INTEGRAL* data) or, equivalently, highest inactivity duty cy-

cle of $67.2^{+4.9}_{-5.7}\%$ (according to the definition¹ of [Romano et al. 2009b, 2014a](#) based on Swift/X-Ray Telescope (XRT) data). IGR J08408–4503 was discovered on 2006 May 15 with the International Gamma-Ray Astronomy Laboratory (*INTEGRAL*) ([Götz et al. 2006](#)). This system hosts a compact object orbiting around HD 74194, a O8.5 Ib-II(f)p supergiant ([Sota et al. 2014](#)), has an orbital period of $P = 9.5436 \pm 0.0002$ d, and a high eccentricity of $e = 0.63 \pm 0.03$ ([Gamen et al. 2015](#)). Its distance from the Sun, derived from the parallax measurements in the *Gaia* Data Release 2, is $d = 2.21^{+0.14}_{-0.16}$ kpc ([Bailer-Jones et al. 2018](#)). IGR J08408–4503 shows, similarly to other SFXTs, sporadic X-ray outbursts composed by flares reaching X-ray luminosities of $L_x \approx 10^{36}$ erg s⁻¹ (e.g. [Romano et al. 2009a](#)). During the quiescence, its X-ray luminosity is of the order of $L_x \approx 10^{32-33}$ erg s⁻¹ ([Sidoli et al. 2010](#); [Bozzo et al. 2010](#)). The X-ray spectrum in the energy range 0.5–50 keV during the flares can be described by an absorbed power law with high-energy cut-off and a black-body component ([Sidoli et al. 2009](#), [Romano et al. 2009a](#), [Götz et al. 2007](#), [Leyder et al. 2007](#)). No cyclotron resonance scattering features have been detected at hard energies. In quiescence, the 0.3–10 keV spectrum is well described by a soft thermal plasma model and a hard power law, with two different absorptions, indicating that IGR J08408–4503 also accretes at low luminosity ([Sidoli et al. 2010](#); [Bozzo et al. 2010](#)). Mechanisms involving accretion of inhomogeneous stellar winds ([Romano et al. 2009a](#)), and gating mechanisms ([Götz et al. 2007](#); [Bozzo et al. 2010](#))

¹ The duty cycle of inactivity is the time each source spends undetected down to a flux limit of $1-3 \times 10^{-12}$ erg cm⁻² s⁻¹, the limiting flux for a 1 ks XRT exposure.

Table 1. Bright X-ray flares ($L_x \gtrsim 10^{36}$ erg s $^{-1}$) detected by *INTEGRAL*, *Swift*/BAT, *MAXI*, and long *Suzaku* and *XMM-Newton* observations of IGR J08408–4503 at lower luminosities.

Time (MJD)	Instrument	Orbital phase	Reference
52821.839	<i>INTEGRAL</i>	0.071	<i>a</i>
53870.770	<i>INTEGRAL</i>	0.980	<i>a</i>
58489.526	<i>INTEGRAL</i>	0.944	this work
53853.223	<i>Swift</i> /BAT	0.141	<i>b</i>
54012.613	<i>Swift</i> /BAT	0.843	<i>b</i>
54214.559	<i>Swift</i> /BAT	0.003	<i>b</i>
54652.887	<i>Swift</i> /BAT	0.932	<i>b</i>
54730.320	<i>Swift</i> /BAT	0.046	<i>b</i>
55071.953	<i>Swift</i> /BAT	0.843	<i>b</i>
55283.664	<i>Swift</i> /BAT	0.026	<i>b</i>
55798.035	<i>Swift</i> /BAT	0.923	<i>b</i>
56475.341	<i>Swift</i> /BAT	0.893	<i>c</i>
55914.029	<i>MAXI</i>	0.077	<i>d</i>
57640.955	<i>MAXI</i>	0.028	<i>d</i>
58451.443	<i>MAXI</i>	0.953	<i>d</i>
55176.746–55179.500	<i>Suzaku</i>	$-0.177 \leq \phi \leq +0.112$	<i>e</i>
54249.391–54249.918	<i>XMM-Newton</i>	$0.653 \leq \phi \leq 0.708$	<i>f</i>

Notes. *a*: Götz et al. (2007); *b*: Romano et al. (2014b); *c*: Romano et al. (2013); *d*: List of the detection of transient sources reported in <http://maxi.riken.jp/alert/novae/index.html>; *e*: Sidoli et al. (2010); *f*: Bozzo et al. (2010).

have been proposed for this source. In light of the orbital parameters recently derived by Gamen et al. (2015) from spectroscopic optical observations, we investigated the possible accretion mechanisms responsible for the X-ray variability exhibited by IGR J08408–4503. For this work we made use of *INTEGRAL*, *Swift*/Burst Alert Telescope (BAT), Monitor of All-sky X-ray Image (*MAXI*), *Suzaku*, and *XMM-Newton* data previously published (Götz et al. 2007; Leyder et al. 2007; Romano et al. 2009a; Sidoli et al. 2009; Bozzo et al. 2010; Sidoli et al. 2010) and of *INTEGRAL* data analysed in this work for the first time (Sect. 2). The results are discussed in Sect. 3.

2. Observations of IGR J08408–4503

2.1. Collection of previous X-ray observations

We calculated the orbital phases and the position along the orbit of all the bright flares ($L_x \gtrsim 10^{36}$ erg s $^{-1}$) of IGR J08408–4503 observed since its discovery by *Swift*/BAT, *INTEGRAL*, and *MAXI*. For *INTEGRAL*, we also searched for previously unreported flares in all the archival data and we found one (see Sect. 2.2). In addition, we considered the long monitoring (about three days) performed by *Suzaku*, where IGR J08408–4503 was observed accreting at an X-ray luminosity of $10^{32} - 10^{33}$ erg s $^{-1}$ (Sidoli et al. 2010), and an *XMM-Newton* observation where IGR J08408–4503 was caught at a similar low luminosity level (Bozzo et al. 2010). We assumed the orbital parameters derived by Gamen et al. (2015), and we also kept their assumption about the mass of the donor star, $M_d = 33 M_\odot$ (based on the calibration work of Martins et al. 2005) and their derivation for the mass of the NS, $M_{\text{ns}} = 1.61 M_\odot$. The times and orbital phases of these observations are shown in Table 1. The positions of these observations along the orbit are shown in Fig. 1.

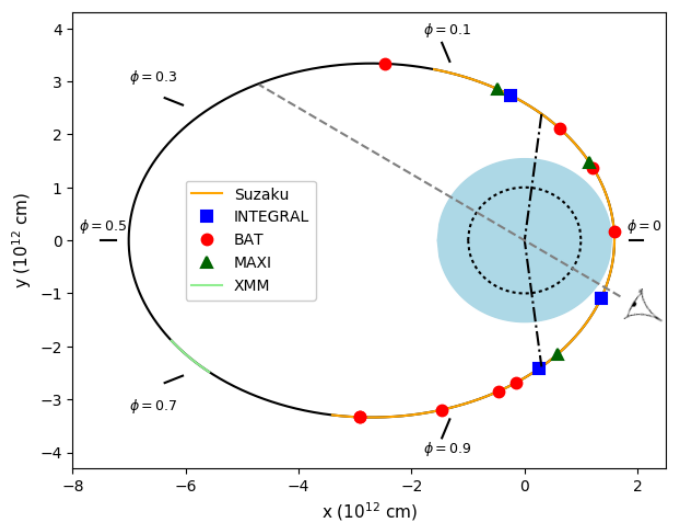


Fig. 1. Positions of the bright X-ray flares ($L_x \gtrsim 10^{36}$ erg s $^{-1}$) and the *Suzaku* and *XMM-Newton* observations along the orbit of IGR J08408–4503. The blue circle shows the size of the donor star ($R_d = 22 R_\odot$), the dotted circle shows the size of the Roche lobe at the periastron, and the dot-dashed line shows the fraction of the orbit around the periastron where the system is in RLO.

2.2. Analysis of *INTEGRAL* data

We analysed all the available archival data collected by the *INTEGRAL* Soft Gamma-Ray Imager (ISGRI; Lebrun et al. 2003) detector of the Imager on Board *INTEGRAL* Satellite (IBIS; Ubertini et al. 2003) on board *INTEGRAL* (Winkler et al. 2003), from 2003 January 11 to 2019 February 17 (corresponding to 3222 pointings). The IBIS instrument operates in the ~ 20 keV–10 MeV energy range. We performed the analysis using the Off-line Science Analysis (OSA) software provided by the ISDC Data Center for Astrophysics². We only considered observations with IGR J08408–4503 within 12° from the centre of the IBIS/ISGRI field of view (FOV). In addition to the previously known flares, IGR J08408–4503 showed a bright flare in 2019 January 6 at 12:11:31.200 UTC (see Fig. 2, top panel). During the flare ($T_{\text{start}} = 58489.508$ MJD; $T_{\text{stop}} = 58489.541$ MJD) the spectrum can be described by a power law with high energy cut-off, with photon index of 2.5 ± 0.2 , $E_c = 54.0^{+8.9}_{-8.8}$ keV, $E_f = 34^{+31}_{-16}$ keV ($\chi^2_{\text{red}} = 1.36$, 7 d.o.f.) Uncertainties are given at 68% confidence level. We added systematic uncertainties of 3% to the IBIS/ISGRI spectrum. The peak flux is $\sim 2.7 \times 10^{-9}$ erg cm $^{-2}$ s $^{-1}$ (30–80 keV), corresponding to a luminosity of $L_x \approx 1.6 \times 10^{36}$ erg s $^{-1}$ (Fig. 2, bottom panel). We searched for periodicities using IBIS/ISGRI data during the flare (i.e. we defined a good time interval with $T_{\text{start}} - T_{\text{stop}}$ defined above) by selecting events with pixel fraction illuminated by IGR J08408–4503 equal to one from the detector shadowgrams obtained before the reconstruction (deconvolution) of the sky image. In addition, we also performed this search using binned light curves (bin= 0.1 s) obtained with the OSA binning tool *ii_light*. In both cases the Solar System barycentre correction was applied. The search for periodic signals was performed in the range 0.01–1000 s with the event files and 0.2–1000 s with the binned

² OSA v.11 for data collected after *INTEGRAL* revolution 1626: https://www.isdc.unige.ch/integral/download/osa/doc/11.0/osa_um_ibis/man.html, and OSA v.10.2 otherwise: https://www.isdc.unige.ch/integral/download/osa/doc/10.2/osa_um_ibis/man.html

light curve. We used the Z^2 and the Lomb-Scargle periodogram techniques (Scargle 1982; Lomb 1976; Bucchieri et al. 1983). No statistically significant periodicity was detected. During this flare, IGR J08408–4503 was outside of the JEM-X FOV. We estimated a 2σ upper limit to the pulsed fraction of 55% and 50% for periods in the ranges 0.2–10 s and 10–1000 s, respectively.

We performed an analysis of all the available photometric data of IGR J08408–4503 collected by the Optical Monitoring Camera (OMC; Mas-Hesse et al. 2003) on board *INTEGRAL*. This camera has a much smaller FOV compared to ISGRI ($4.979^\circ \times 4.979^\circ$) and it is equipped with a Johnson *V* filter. We reduced the data following the prescriptions of the OMC Analysis User Manual³. Since IGR J08408–4503 is bright ($V \approx 7.5$), we considered only the observations with integration time of 10 s because they allow a good measurement of the magnitude without pixel saturation. We thus selected 124 “good” observations of IGR J08408–4503. Figure 3 shows the light curve as a function of time (top panel) and folded with the orbital period of the source (bottom panel), where phase zero corresponds to the periastron passage. The OMC light curve presented here is, to the best of our knowledge, the most detailed and long photometric monitoring of this source. It shows irregular variability with amplitudes up to ~ 0.14 .

3. Results and discussion

The bright X-ray flares of IGR J08408–4503 previously observed and the last flare detected by *INTEGRAL* and reported in this work are clustered around the periastron, within the orbital interval $\sim 0.84 - 0.14$. The long *Suzaku* observation carried out during the same orbital phases of the flares shows a variable low luminosity state of $\sim 10^{32-33} \text{ erg s}^{-1}$ (Sidoli et al. 2010), which implies a luminosity dynamic range of $\Gamma = 10^3 - 10^4$. Based on the calibration of O star parameters provided by Martins et al. (2005), we found that the average radius of the donor star of IGR J08408–4503, $R_d \approx 22 R_\odot$, is larger than the size of its Roche lobe at the periastron, $R_{1,d} \approx 14.4 R_\odot$, calculated with the formula provided by Eggleton (1983, equation 2) assuming the stellar and orbital parameters previously adopted in this work (see Fig. 1, where the dotted circle shows the $R_{1,d}$ at periastron). The dot-dashed line of Fig. 1 shows the fraction of the orbit where the system is in Roche lobe overflow (RLO; i.e. when $R_d \geq R_{1,d}$). We obtained similar results using the calibration of O star parameters provided by Vacca et al. (1996), although in this case the radius of the donor star results to be larger than the periastron distance. The strong tidal interaction that would be caused by the periodic passage of the NS inside the photosphere of the companion star would likely produce a displacement of a significant fraction of the photosphere itself. Consequently, this would lead to a periodic optical luminosity and column density variability. In addition, the high mass accretion rate at $\phi = 0$, where the NS would be inside the companion star, would cause a higher rate of X-ray flares at $\phi \approx 0$. Since none of these properties has been observed in IGR J08408–4503, this suggests that the radius of the donor star is likely to be smaller than the periastron distance. With the information about the visual magnitude measured by the OMC ($m_v = 7.55 \pm 0.05$), the distance ($d = 2.21^{+0.14}_{-0.16}$ kpc), the effective temperature ($T_{\text{eff}} = 32274 \pm 1800$ K, Martins et al.

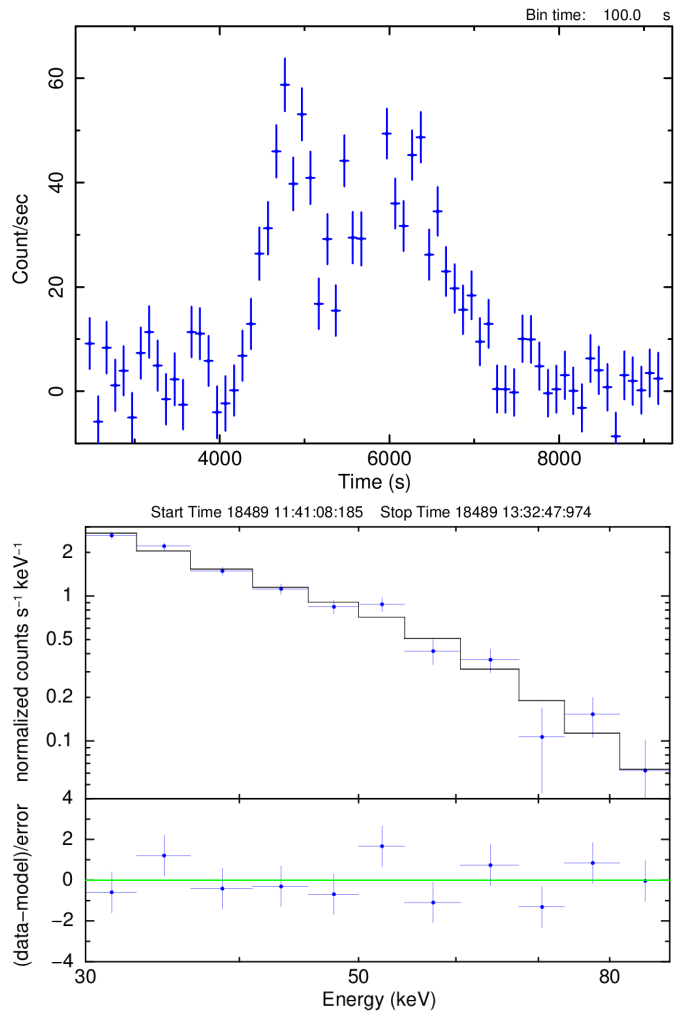


Fig. 2. *Top panel:* 30–50 keV light curve of the 2019 January flare of IGR J08408–4503. *Bottom panel:* IBIS/ISGRI spectrum of the flare of IGR J08408–4503.

2005)⁴, the absorption ($N_{\text{H}} = 5.5 \pm 0.4 \times 10^{21} \text{ cm}^{-2}$)⁵, and assuming that the radiation from the donor star is a black body, we found a radius of $R_d = 23.4^{+7.0}_{-5.4} R_\odot$. Although it is not well constrained, we note that it is consistent with the values of R_d mentioned above for this type of stars, and thus consistent with the RLO hypothesis. Using the *B* and *R* magnitudes from the United States Naval Observatory B1 (USNO B1) catalogue (Monet et al. 2003) and the extinction law of Cardelli et al. (1989), we obtained $R_d = 23.3^{+10.1}_{-8.3} R_\odot$. The larger uncertainty of this estimate is due to the worst photometric accuracy of the USNO B1 catalogue with respect to the OMC measurements.

Regös et al. (2005) studied the mass transfer in eccentric binaries in which RLO is expected to occur. They found that for high eccentricities ($e \geq 0.6$), a high mass transfer at pe-

⁴ The error on T_{eff} is a rough estimate obtained by the values in tables 3 and 6 of Martins et al. (2005).

⁵ Obtained from the *XMM-Newton* observations of IGR J08408–4503 during the low luminosity state and reported in Sidoli et al. (2013). We only considered the absorption affecting the MEKAL (Mewe-Kaastra-Liedahl) component to exclude the absorption produced around the accreting compact object. The column density is obtained from the weighted mean of the values reported in table 3 of Sidoli et al. (2013), with uncertainties corrected to 1σ confidence level. We converted N_{H} to A_{V} using the formula presented in Güver & Özel (2009) and taking into account the uncertainties reported in their paper.

³ https://www.isdc.unige.ch/integral/download/osa/doc/11.0/osa_um_omc.pdf

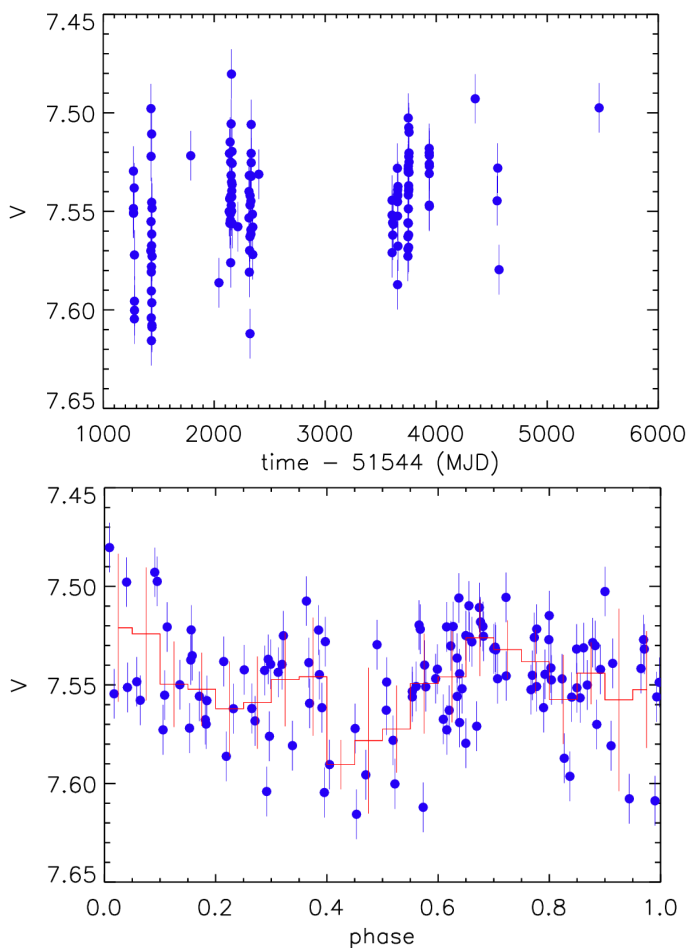


Fig. 3. *Top panel:* V band light curve of IGR J08408–4503, observed by the OMC. *Bottom panel:* The same light curve, folded with the orbital period found by [Gamen et al. \(2015\)](#). Phase zero corresponds to the periastron passage. Blue points have a time duration of 10 s. Red line shows the binned light curve. The error bars of the red line are the standard deviation obtained from the magnitude values of the blue points and show their dispersion.

riastron is possible, and matter leaves the donor star through the Lagrangian points L_1 and L_2 . A fraction of the matter leaving L_2 falls back to the star, while the other part does not remain bound to the star. Similarly, part of the stream passing through L_1 leaves the system, while most of it forms an accretion disc. Given the orbital parameters, this mechanism might be at work in IGR J08408–4503. The stream of matter lost by the donor star through L_2 could make the environment around the binary system not “cleaned”: slow and denser regions of gas could be crossed by the NS along its orbit, feeding further the NS. In this scenario, the aperiodic optical variability shown by IGR J08408–4503 in the OMC light curve (Fig. 3) might be caused by sporadic absorption of dense clumps of gas lost from L_1 and L_2 crossing the line of sight, or by tidal interactions between the two stars, which disrupts the outer layers of the photosphere of HD 74194. As previously mentioned, another possible peculiarity of IGR J08408–4503 is the short distance between the NS and the surface of the donor star at periastron, that is just $\approx 6 \times 10^{10}$ cm, assuming $R_d = 22 R_\odot$ (the periastron distance is $\sim 22.9 R_\odot$). This would imply that there is likely not enough space for an accretion disc which, once formed during the orbital phases preceding $\phi \approx 0$, might be disrupted because of the tidal

interactions with the donor star. A slightly smaller radius for the supergiant would increase the chances of survival of the disc; for example, if $R_d \approx 19 R_\odot$, i.e. the radius of HD 75211, a O8.5 II(f) star ([Markova et al. 2018](#)).

If the NS of IGR J08408–4503 is fed by an accretion disc, the high X-ray variability could be caused by transitions from inhibition of accretion (during low luminosity states) to the accretion regime (during flares). This would be possible for small variations of the mass rate gravitationally captured by the NS (\dot{M}_c) if the corotation radius

$$r_{\text{co}} = [GM_{\text{ns}}P_{\text{spin}}^2/(4\pi^2)]^{1/3}, \quad (1)$$

defined as the distance from the NS where material in a Keplerian orbit corotates with the NS, is comparable in size to that of the inner radius of the disc; this is defined as the distance from the NS where the magnetic pressure and pressure exerted by the matter in the disc are balanced:

$$r_{\text{in}} \approx \xi[\mu^4/(GM_{\text{ns}}\dot{M}_c^2)]^{1/7}. \quad (2)$$

In these equations, M_{ns} is the mass of the NS and P_{spin} its spin period, $\xi \approx 0.4$ in case of accretion from a Keplerian disc or 1 in case of spherical accretion (see [Revnivtsev & Mereghetti 2015](#) and references therein), μ is the magnetic dipole moment, $B = 2\mu/R_{\text{ns}}^3$ is the magnetic field strength at the poles, and R_{ns} is the radius of the NS.

If $r_{\text{in}} > r_{\text{co}}$, the centrifugal barrier sets in. During this stage the accretion is expected to be negligible. Once $r_{\text{in}} \approx r_{\text{co}}$, the centrifugal barrier is overcome and accretion sets in. In this framework, we can consider the following very simplified calculations. We assume that during the accretion state the X-ray luminosity is given by $L_{\text{acc}} \approx GM_{\text{ns}}\dot{M}_c/R_{\text{ns}}$. When the magnetic barrier activates ($r_{\text{in}} = r_{\text{co}}$) for small variations of \dot{M}_c , which for simplicity can be considered constant in our calculations, this value reduces to $L_{\text{gate}} \approx GM_{\text{ns}}\dot{M}_c/r_{\text{co}}$. Therefore, r_{co} can be expressed as a function of $\Gamma = L_{\text{acc}}/L_{\text{gate}}$,

$$r_{\text{co}} = R_{\text{ns}}\Gamma. \quad (3)$$

Assuming again $r_{\text{in}} = r_{\text{co}}$, from Eq. 3 and Eq. 2, we obtain the following relation between μ , L_{acc} , and Γ :

$$\mu = \xi^{-7/4}R_{\text{ns}}^{9/4}\Gamma^{7/4}G^{-1/4}M_{\text{ns}}^{-1/4}L_{\text{acc}}^{1/2}. \quad (4)$$

Assuming $\Gamma = 10^3 - 10^4$, $L_{\text{acc}} = 10^{36}$ erg s $^{-1}$, $R_{\text{ns}} = 12$ km, $M_{\text{ns}} = 1.61 M_\odot$, and $\xi = 0.4$, we find $\mu = 1.1 \times 10^{31} - 6.2 \times 10^{32}$ G cm 3 and hence $B = 1.3 \times 10^{13} - 7.2 \times 10^{14}$ G. From Eq. 1 we then obtain $P_{\text{spin}} = 18 - 570$ s, while the corotation radius is $r_{\text{co}} = 1.2 \times 10^9 - 1.2 \times 10^{10}$ cm. These calculations indicate a magnetar nature for the compact object of IGR J08408–4503 for a wide range of values of Γ . The lack of clear signatures for the presence of an accretion disc in the X-ray spectrum of IGR J08408–4503 in the spectral analyses presented in previous works is in agreement with the high magnetic field and low spin period inferred from the calculations presented above. If $r_{\text{in}} = r_{\text{co}}$, from the equation of the spectrum emitted by an optically thick accretion disc (e.g. [Frank et al. 2002](#), equation 5.43),

$$T(r) = \left[\frac{3GM\dot{M}_c}{8\pi r^3\sigma} \left(1 - \sqrt{\frac{R_{\text{ns}}}{r}} \right) \right]^{1/4},$$

where σ is the Stefan-Boltzmann constant, we obtain $T(r_{\text{in}}) \sim 1$ eV, which means that an accretion disc would be too cold to be observed in X-ray.

A possible problem for the accretion disc scenario is that the bright flares ($L_x \gtrsim 10^{36}$ erg s $^{-1}$) are clustered in an interval of the orbit around the periastron ($\Delta\phi \approx 0.3$). We expect a higher mass accretion rate in this interval of the orbit because of RLO and because the stellar wind is denser and slower closer to the donor star. However if the accretion is mediated by a disc, the NS would feel effects of the mass accretion rate variability with a time delay, which is of the order of the viscous timescale (\sim days; for IGR J08408–4503) and not necessarily in phase with the orbital period. One possibility previously mentioned is that the size of the donor star is so large that the accretion disc is disrupted by tidal interactions at $\phi \approx 0$. In this case, the accretion is spherically symmetric along the orbit—no delay effects are expected—and the gating mechanism model of [Bozzo et al. \(2008\)](#) can be applied. In this framework, the denser streams of gas expelled by the donor star from L_2 and L_1 , and sporadically accreted by the NS during the periastron passages, could trigger the transitions from low to higher luminosity accretion regimes described in [Bozzo et al. \(2008\)](#). Instead, if the size of the donor star is such that the accretion disc survives at $\phi \approx 0$, a two-stream accretion mechanism similar to that proposed by [Lipunov \(1980, 1982\)](#) could be in operation. During RLO, an accretion disc forms. In addition, the NS accretes from the strong stellar wind produced by the companion star all along the orbit. [Lipunov \(1982\)](#) pointed out that during the propeller state, which could be the low luminosity state of IGR J08408–4503, the spherical stream collides with the magnetosphere and slowly settles down on the NS equatorial plane to form or feed an already existing accretion disc. In this case, the clustering of the bright flares around the periastron might be caused by the sudden overcoming of the centrifugal barrier ($r_{\text{in}} \lesssim r_{\text{co}}$) triggered by instabilities produced by the enhanced accretion produced by the spherical stream around the periastron. The interactions between the two accretion streams are still poorly studied. [Mustsevoi & Khoperskov \(1991\)](#) pointed out that they could lead to instabilities in the accretion disc.

IGR J08408–4503 could not be an isolated case of accretion through RLO among SFXTs. To the best of our knowledge there are at least two other cases: IGR J16418–4532 and IGR J16479–4514. Their orbital, stellar, and Roche lobe properties are reported in [Table 2](#). While for reasonable values of the radius, the donor star of IGR J16479–4514 could still underfill its Roche lobe (as pointed out by [Sidoli et al. 2013](#)), for IGR J16418–4532 the radius of the donor star is clearly too large compared to its Roche lobe (see [Sidoli et al. 2012](#) and [Drave et al. 2013](#) for further discussions). In particular, we note that with the parameters reported in [Table 2](#), the radius of the donor star of IGR J16418–4532 would be slightly larger than the orbital size, which is obviously not possible. According to [Crowther et al. \(2006\)](#), the radii of B 0.5Ia stars can be significantly different between different stars. Although on average they are ~ 33 – $34 R_{\odot}$, for example another star of similar spectral type, κ Ori, has $R_* \approx 22.2 R_{\odot}$, which would still imply $R_* \gtrsim R_{1,d}$ for IGR J16418–4532.

The formation of accretion discs through RLO could therefore be a property common to a group of SFXTs. This hypothesis, however, requires further investigation. Therefore, in-depth optical observations and X-ray studies aimed at determining more precisely the masses and radii of the SFXTs reported in [Table 2](#) and the presence of an accretion disc are of fundamental importance.

Acknowledgements. We thank the anonymous referee for constructive comments that helped to improve the paper. This paper is based on data from observations with INTEGRAL satellite, an ESA project with instruments and science data

Table 2. List of SFXTs which might experience RLO in a fraction or along the entire NS orbit. The quantity $R_{1,d}$ is the size of the Roche lobe of the donor star during the periastron passage.

Source name	Sp. Type	Mass (M_{\odot})	Radius (R_{\odot})	P_{orb} (d)	$R_{1,d}$ (R_{\odot})
IGR J08408–4503	O8.5 Ib-II(f)p	33 ¹	22 ¹	9.4536 ²	14.38
IGR J16418–4532	BN 0.5Ia ³	33 ⁴	33.8 ⁴	3.7389 ⁵	21.14
IGR J16479–4514	O9.5 Iab ⁶	27.8 ¹	22 ¹	3.3193 ^{7,8}	18.12

Notes. ¹: [Martins et al. \(2005\)](#); ²: [Gamen et al. \(2015\)](#); ³: [Coleiro et al. \(2013\)](#); ⁴: [Searle et al. \(2008\)](#); ⁵: [Levine et al. \(2011\)](#); ⁶: [Nespoli et al. \(2008\)](#); ⁷: [Jain et al. \(2009\)](#); ⁸: [Romano et al. \(2009a\)](#). We assumed $e = 0.63$ for IGR J08408–4503 and $e = 0$ for the other SFXTs because of the lack of an orbital modulation of the X-ray flux ([Sidoli et al. 2012](#); [Drave et al. 2013](#)).

centre funded by ESA member states (especially the PI countries: Denmark, France, Germany, Italy, Spain, and Switzerland), Czech Republic and Poland, and with the participation of Russia and the USA. This work is supported by the Bundesministerium für Wirtschaft und Technologie through the Deutsches Zentrum für Luft und Raumfahrt (grant FKZ 50 OG 1602). P.R. acknowledges contract ASI-INAF I/004/11/0. We acknowledge financial contribution from the agreement ASI-INAF n.2017-14-H.O.

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